

Deriving precise parameters for cool solar-type stars

Optimizing the iron line list[★]

M. Tsantaki^{1,2}, S. G. Sousa^{1,3}, V. Zh. Adibekyan¹, N. C. Santos^{1,2}, A. Mortier^{1,2}, and G. Israelian³

¹ Centro de Astrofísica, Universidade do Porto, Rua das Estrelas, 4150-762 Porto, Portugal
e-mail: Maria.Tsantaki@astro.up.pt

² Departamento de Física e Astronomia, Faculdade de Ciências, Universidade do Porto, Rua do Campo Alegre, 4169-007 Porto, Portugal

³ Instituto de Astrofísica de Canarias, E-38200 La Laguna, Tenerife, Spain

Received XXXX; accepted XXXX

ABSTRACT

Context. Temperature, surface gravity, and metallicity are basic stellar atmospheric parameters necessary to characterize a star. There are several methods to derive these parameters and a comparison of their results often shows considerable discrepancies, even in the restricted group of solar-type FGK dwarfs.

Aims. We want to check the differences in temperature between the standard spectroscopic technique based on iron lines and the Infrared Flux Method (IRFM). We aim to improve the description of the spectroscopic temperatures especially for the cooler stars where the differences between the two methods are higher, as presented in previous work.

Methods. Our spectroscopic analysis is based on the iron excitation and ionization balance, assuming Kurucz model atmospheres in LTE. The abundance analysis is determined using the code MOOG. We optimize the line list using a cool star (HD 21749) with high resolution and high signal-to-noise spectrum, as a reference in order to check for weak, isolated lines.

Results. We test the quality of the new line list by re-deriving stellar parameters for 451 stars with high resolution and signal-to-noise HARPS spectra, that were analyzed in a previous work with a larger line list. The comparison in temperatures between this work and the latest IRFM for the stars in common shows that the differences for the cooler stars are significantly smaller and more homogeneously distributed than in previous studies for stars with temperatures below 5000 K. Moreover, a comparison is presented between interferometric temperatures with our results that shows good agreement, even though the sample is small and the errors of the mean differences are large. We use the new line list to re-derive parameters for some of the cooler stars that host planets. Finally, we present the impact of the new temperatures on the [Cr I/Cr II] and [Ti I/Ti II] abundance ratios that previously showed systematic trends with temperature. We show that the slopes of these trends for the cooler stars become drastically smaller.

Key words. techniques: spectroscopic – stars: fundamental parameters

1. Introduction

Temperature (T_{eff}), surface gravity ($\log g$), and metallicity ($[Fe/H]$, where iron is used as a proxy) are basic atmospheric parameters necessary to characterize a star, as well as to determine other indirect and fundamental parameters, such as mass, radius and age that are acquired in combination with stellar evolutionary models (Girardi et al. 2000). Precise and accurate stellar parameters are also essential in exoplanet searches. The light curve of a transiting planet orbiting a star gives information of the planetary radius always in dependence of the stellar radius ($R_p \propto R_*$). Moreover, the mass of the planet, or the minimum mass in case the inclination of the orbit is not known, is determined from the radial velocity technique only if the mass of the star is known ($M_p \propto M_*^{2/3}$). Therefore, the determination of the planetary radius and mass requires a combination of transiting and radial velocity data (Ammler-von Eiff et al. 2009; Torres et al. 2008, 2012). Apart from the derivation of planetary properties (mass, radius, density), stellar parameters can be used to reveal correlations between planets and their host stars that

will give insights on their formation and evolution mechanisms. Several correlations have been studied so far with very interesting results, such as stellar metallicity and planet frequency (e.g. Santos et al. 2004; Udry et al. 2007; Sousa et al. 2011), stellar metallicity and planetary mass (Guillot et al. 2006), stellar metallicity and planetary orbital periods (Sozzetti 2004), stellar temperature and obliquities (Albrecht et al. 2012; Winn et al. 2010), metallicity and planet radius (Buchhave et al. 2012). It is obvious that in order to correctly characterize the planets and furthermore, to statistically verify such correlations, well-determined stellar parameters are required.

Hundreds of stellar spectra are available from radial velocity planet search programs (Udry et al. 2000; Vogt et al. 2000; Mayor et al. 2003; Bouchy et al. 2009; Lo Curto et al. 2010). These spectra are obtained with high resolution spectrographs and a combination of them gives high signal-to-noise (S/N) that makes the spectroscopic analysis more powerful.

The study of the $H\alpha$ wings (Fuhrmann 2004), the excitation and ionization balance of iron lines (Santos et al. 2004), spectral synthesis techniques (Valenti & Fischer 2005), line ratios (or line depth ratios) (Gray 1994) are the basic spectroscopic techniques that along with fundamental techniques such as photometry and interferometry can be used for the determination of the effective

[★] Based on observations collected at the La Silla Paranal Observatory, ESO (Chile) with the HARPS spectrograph at the 3.6-m telescope (ESO runs ID 072.C-0488).

temperature. A comparison between these different methods can show considerable discrepancies in their results. Even in the restricted group of solar-type stars, the effective temperatures obtained with these methods can differ significantly (e.g. Kovtyukh et al. 2003; Ramírez & Meléndez 2004; Casagrande et al. 2006; Sousa et al. 2008).

The temperature determination becomes more difficult when we focus on K-type stars. The difficulties in these stars with $T_{\text{eff}} \lesssim 5000$ K emerge from their line crowded spectra that cause strong blending. Blending can be a considerable problem if one uses the standard technique based on the iron equivalent widths (EWs). Lines cannot be easily resolved and the continuum placement becomes more difficult, causing bad measurement of the EWs and hence, makes the calculation of stellar parameters ambiguous. Therefore, it is important to select carefully the iron lines in such manner that will eliminate the blending effects, especially for cool stars. In addition, the choice of the atomic parameters influences the abundance determination. Some authors calculate the atomic parameters using the Sun as a reference to avoid the errors that emerge from the theoretical or laboratory values. In that way, the atomic parameters for stars that are different from the Sun, i.e. too hot or too cool can be no longer accurate enough.

The accuracy in the fundamental parameters of planet host stars is of great importance to planetary studies since these errors can propagate to the planetary properties (e.g. Pepe et al. 2011). In addition, stellar parameters affect the determination of other element abundances ($[X/H]$) and can introduce potential biases in their abundance calculations (e.g. Neves et al. 2009; Adibekyan et al. 2012b).

In this paper, we use the data from previous work that performed the standard spectroscopic analysis for a sample of solar-type stars presented in the paper of Sousa et al. (2008) (hereafter SO08). This sample is part of the High Accuracy Radial velocity Planet Searcher (HARPS) Guaranteed Time Observations (GTO) survey that is composed of slow rotators and low activity FGK stars in order to detect low mass planets.

A comparison of these spectroscopic results with the Infrared Flux Method (IRFM) indicates a disagreement between the effective temperatures only for the cooler stars of the sample with temperatures below ~ 5000 K. Motivated by that, we compile an optimized line list to improve the accuracy of the stellar parameters for the cooler stars and compare our results with other independent methods (IRFM, interferometry). This paper is organized as follows: in Sect. 2 we describe the stellar sample, the spectroscopic method and the criteria to select an optimized line list. We compare our results with the previous work of SO08 in Sect. 3. In Sect. 4, we compare the spectroscopic surface gravities with the trigonometric ones and address the ionization problem for the cooler stars of our sample. In Sect. 5, we compare the new derived parameters with the IRFM and interferometry. In Sect. 6, we calculate new parameters for some cool planet hosts. Finally, in Sect. 7, we use the new parameters for the cooler stars to reduce the trends of other elements ($[\text{Cr I}/\text{Cr II}]$ and $[\text{Ti I}/\text{Ti II}]$ abundance ratios) with temperature, as reported in Adibekyan et al. (2012b).

2. Stellar sample and previous spectroscopic analysis

The stellar sample, presented in SO08, is composed of 451 stars as part of the HARPS high-precision GTO program at the ESO La Silla 3.6m telescope with the objective to detect low-mass

Table 1. Characteristics of the sample and the reference star, described in SO08.

	T_{eff} (K)	$\log g$ (dex)	$[Fe/H]$ (dex)	$Mass$ M_{\odot}
Lowest	4556	3.68	-0.84	0.37
Highest	6403	4.62	0.39	1.42
Reference star	4723	4.40	-0.02	0.76

extra-solar planets with high radial velocity accuracy (Mayor et al. 2003). It is mainly comprised of dwarf FGK stars selected from a volume-limited sample of the CORALIE survey (Udry et al. 2000). Planet hosts from the southern hemisphere were also added to this sample, forming in total a sample of 451 stars. These stars are slowly-rotating, non-evolved, and low-activity stars, with apparent magnitudes that range from 3.5 to 10.2 and have distances of less than 56 parsec. The spectra have a resolution of $R \sim 110,000$ and 90% of the combined spectra have S/N higher than 200. We point to SO08 for more details.

For this sample, SO08 derived stellar parameters by imposing excitation and ionization equilibrium, based on the measurements of weak iron lines. This method is very effective for FGK stars due to the numerous iron lines in their spectra. Iron abundance is used as a proxy for the overall stellar metallicity.

The line list for their spectroscopic analysis, was composed of 263 Fe I and 36 Fe II lines. The EWs of the lines were measured automatically for all stars using the ARES¹ code (Automatic Routine for line Equivalent widths in stellar Spectra; Sousa et al. 2007). The atomic parameters of the iron lines, namely the oscillator strength values ($\log gf$), were computed by an inverted solar analysis, using a solar model with $T_{\text{eff}} = 5777$ K, $\log g = 4.44$ dex, $\xi_t = 1.0$ m s⁻¹, $\log_{\epsilon}(Fe) = 7.47$ dex.

The spectroscopic analysis was completed assuming Local Thermodynamic Equilibrium (LTE), and using the 2002 version of the abundance determination code MOOG² (Snedden 1973) and a grid of Kurucz Atlas 9 plane-parallel, 1D static model atmospheres (Kurucz 1993).

The best parameters were obtained when the Fe I abundance shows no dependence on the excitation potential, χ_i , and on the reduced equivalent width, $\log W_{\lambda}/\lambda$. Additionally, the mean abundances given by Fe I and the Fe II must be the same (ionization balance) and consistent with those of the input model atmosphere. The input parameters converge to the true ones with an iterative minimization code based on the Downhill Simplex Method (Press et al. 1992), making the total procedure automatic (Santos et al. 2004). Some characteristics of the sample are depicted in Table 1, as described in SO08.

There are different sources of uncertainties that occur in the stellar parameter determination using this method. These errors can be attributed to the uncertainties of the measurements of the EWs, the uncertainties in the atomic parameters and the uncertainties that are intrinsic to the method of ionization and excitation equilibrium. In addition, systematic errors can arise due to the assumptions of the method, such as 1D static atmospheres, NLTE effects (Mashonkina et al. 2011; Bergemann et al. 2012). However, departures from LTE for Fe lines do not need to be

¹ The ARES code is an open source code and can be found at <http://www.astro.up.pt/~sousasag/ares>

² The MOOG code can be downloaded free at <http://verdi.as.utexas.edu/moog.html>

considered for near solar metallicity dwarfs but should be taken into consideration for more evolved or very metal-poor stars (Lind et al. 2012; Ruchti et al. 2013) that are not part of this sample.

Errors in the measurements of the EWs can be minimized by using high quality spectra. In low S/N spectra, weak lines cannot be distinguished from noise and strong lines can be underestimated due to the miscalculation of their wings. The high resolution and high S/N spectra used for this sample, are the best solution to deal with such errors. Since in our spectroscopic analysis the atomic data ($\log gf$) are derived with respect to the Sun, we expect small errors for solar analogs and more significant for cooler and hotter stars.

The errors in the stellar parameters for this analysis are calculated by varying each parameter (temperature, surface gravity and microturbulence) by a standard value. The uncertainty in ξ_t is determined from the standard deviation in the slope of the least-squares fit of $\log_\epsilon(\text{Fe I})$ versus $\log W_\lambda/\lambda$. The uncertainty in T_{eff} is determined from the uncertainty in the slope of the least-squares fit of $\log_\epsilon(\text{Fe I})$ versus χ_I in addition to the uncertainty in the slope due to the uncertainty in ξ_t . The uncertainty in $\log g$ is derived from the contribution of the uncertainty in T_{eff} and microturbulence in addition to the scatter error of the Fe II abundance (measured as σ/\sqrt{N} , σ is the standard deviation and N the number of lines). The uncertainty in the $\log_\epsilon(\text{Fe I})$ is the sum of the squared uncertainties due to the error in T_{eff} , ξ_t , and the scatter error of Fe I abundance. The use of many iron lines can reduce this type of uncertainty, assuming that the majority of the lines are independent and of good quality.

2.1. Building a stable line list for the cooler stars

A reliable line list is comprised of lines that can be accurately measured, which usually means unblended lines. In addition, lines must be unsaturated, cover a wide range in excitation potential and have accurate atomic data. Temperature, as well as the other stellar parameters, is strongly correlated with the equivalent width. This sensitivity emerges from the excitation and ionization processes that follow the exponential and power dependencies with temperature that are defined by the well-known Boltzmann and Saha equations.

For the cooler stars line blending is severe, which makes the measurements of the EWs problematic. In particular, blending effects cause an overestimation of the EWs as two blended lines cannot be resolved.

Another bias in the EW measurements may come from the fitting of strong lines. Gaussian fitting is a good approximation for weak lines and it can be reliable up to 200 mÅ based on our experience, whereas a Voigt profile should be used for stronger lines. Saturated lines that deviate significantly from the linear part of the curve of growth should also be avoided in the abundance analysis. The EW predicted by the models of strong lines that are highly saturated, is quite dependent on microturbulence. A wrong estimation of microturbulence, will then produce errors in the abundance of any highly saturated line.

On the other hand, weak lines that are strongly blended could lead to a sub-estimation of the continuum and consequently of the EW. This effect, however, is less significant. An overestimation in the EWs due to blending, as well as the underestimation of very strong lines could be the reason for the systematic raise in temperature that is observed for the cooler stars of SO08. In addition, the reduced equivalent width could also be affected by

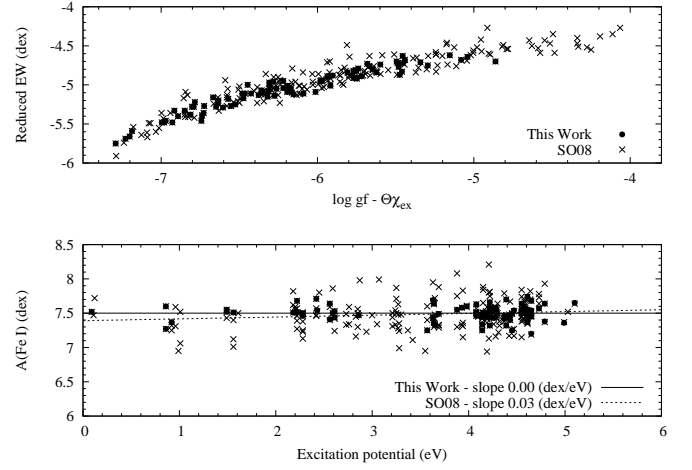


Fig. 1. Upper panel: Curve of growth for both line lists for the reference star computed for the temperature ($\Theta=5040/T_{\text{eff}}$) of this work. Circles represent the reduced EW of the line list of this work and crosses the line list of SO08. Lower panel: The Fe I abundances of the reference star versus the excitation potential. The dashed line shows the positive slope that corresponds to the line list of SO08. The solid line corresponds to the line list of this work and the slope is obviously zero.

Table 2. Sample of the line list used for the spectroscopic analysis with the atomic parameters of Fe I and Fe II as well as the corresponding EWs of the reference star HD 21749.

λ (Å)	χ_I	$\log gf$	Element	EW(mÅ)
4508.28	2.86	-2.403	Fe II	53.0
4520.22	2.81	-2.563	Fe II	72.4
4523.40	3.65	-1.871	Fe I	101.7
4537.67	3.27	-2.870	Fe I	43.2
4551.65	3.94	-1.928	Fe I	41.9
4556.93	3.25	-2.644	Fe I	57.9
4566.52	3.30	-2.156	Fe I	68.6
4574.22	3.21	-2.353	Fe I	55.1
4576.34	2.84	-2.947	Fe II	29.6
...

such biases, leading to correlations with the excitation potential (see Appendix A).

Therefore, our aim is to optimize the iron line list of SO08. With this goal, we use the K-type dwarf HD 21749 with $T_{\text{eff}} = 4723$ K (see Table 1), as reference in order to check for unblended lines in its high S/N spectrum. After visual inspection, we only consider weak, isolated lines that can give good estimation for the local continuum. We avoid strong lines (>200 mÅ) in order to apply Gaussian profiles. For the reference star we show the curve of growth (see Gray 1994) using both line lists (Fig. 1 upper panel). Limiting the EW cut off, we mitigate in large amount the problem of saturated lines and microturbulence. The proof of that mitigation is the fact that the derived temperatures with the new line list agree with other less model dependent methods (see Sect. 5).

Very weak lines (<10 mÅ) were also excluded so that noise is not superposed to these lines. The region of the spectrum below 4500 Å is neglected due to the higher blending. The final

Table 4. Results of the internal comparison for the whole sample and for stars with $T_{\text{eff}} < 5000$ K.

	ΔT_{eff} (K)	$\Delta \log g$ (dex)	$\Delta [\text{Fe}/\text{H}]$ (dex)
This Work – SO08 whole sample	-31 ± 3	-0.023 ± 0.003	0.0002 ± 0.0011
This Work – SO08 $T_{\text{eff}} < 5000$ K	-106 ± 6	-0.07 ± 0.01	0.0127 ± 0.0001

line list³ is compiled with 120 Fe I and 17 Fe II lines, as shown in Table 2.

As mentioned before, the effective temperature is derived when the correlation coefficient between $\log(\text{Fe I})$ and χ_I is zero. In the lower panel of Fig. 1, we demonstrate this correlation for the reference star using the line list of this work and of the work of SO08 using the parameters derived with the line list of this work. The positive slope for the line list of SO08 is translated in an overestimation in temperature of ~ 180 K for this star. In addition, the abundances with the new line list show a smaller scatter which corresponds to smaller errors in the final temperature value.

3. New stellar parameters for 451 FGK stars in the HARPS GTO sample

In order to check the effectiveness of the new line list, we re-derive stellar parameters for the 451 stars of the sample. For consistency, we use the same EWs as in SO08 that were measured automatically for all stars with the ARES code. In addition, we use the same damping parameters and atomic data. After a preliminary determination of the fundamental parameters, we perform a ‘ 3σ clipping’ procedure for lines that contribute with abundances higher than 3σ from the average abundance. This procedure was also applied in SO08.

Table 3 shows the final parameters for a fraction of stars from the sample. Microturbulence (ξ_t) is used as a free parameter and is also derived from this spectroscopic analysis. The correlation of microturbulence with temperature and surface gravity is presented in Appendix B. This calibration can be useful in cases where the value of ξ_t is set fixed. The stellar masses are calculated using the stellar evolutionary models from the Padova group⁴. The errors of the fundamental parameters are internal, attributed to the method. They, thus, represent relative errors and not the absolute accuracy.

3.1. Internal comparison

We present the comparison results with the work of SO08 for temperature, surface gravity and metallicity, respectively in Fig. 2. Temperatures show very good agreement in the high and intermediate ranges. Table 4 shows the mean difference in T_{eff} for the whole sample and for temperatures below 5000 K with their standard errors⁵. The mean difference in T_{eff} is -31 ± 3 ($\sigma=53$) K for the whole temperature range. The significant differences appear, as expected, for stars with temperatures below 5000 K with mean difference $\Delta T_{\text{eff}} = -106 \pm 6$ ($\sigma=54$) K.

³ The full line list is available in electronic form.

⁴ Web interface for stellar mass estimation: <http://stev.oapd.inaf.it/cgi-bin/param>

⁵ The standard errors of the mean (σ_M) are calculated with the following formula: $\sigma_M = \frac{\sigma}{\sqrt{N}}$, σ being the standard deviation.

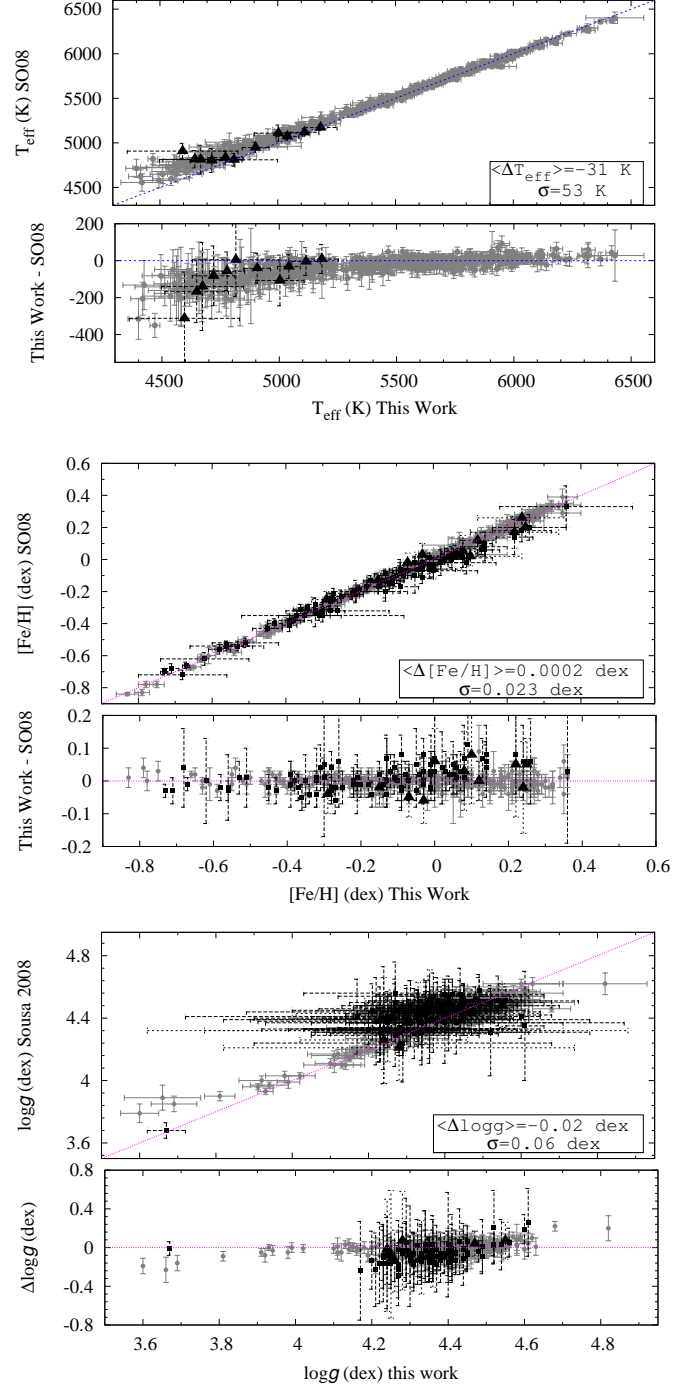
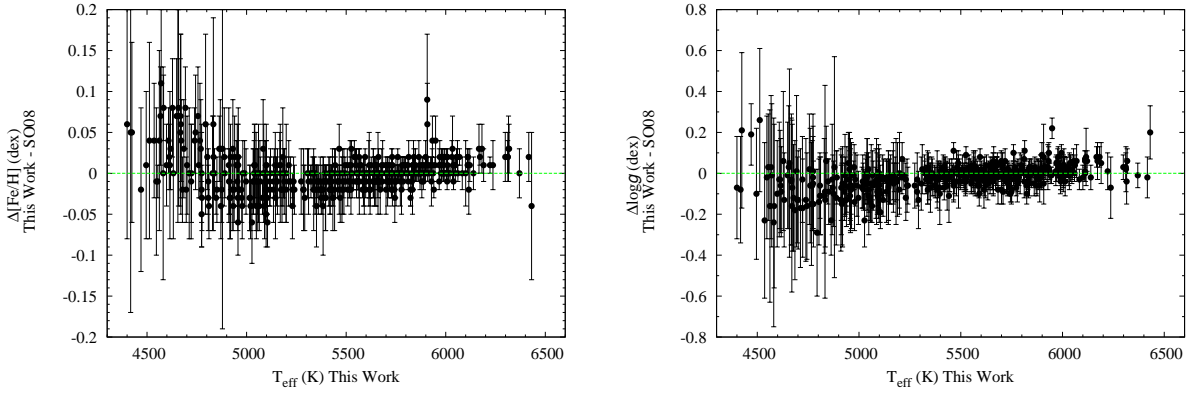


Fig. 2. Comparison between the parameters derived with the cool line list of this work and the results of SO08 for: temperature (top panel), metallicity (middle panel) and surface gravity (bottom panel). ΔT_{eff} corresponds to this work minus SO08. Black squares represent stars with $T_{\text{eff}} < 5000$ K. Triangles represent stars with planets taken from Table 7 (see Sect. 6).

One interesting result is that even though we have considerably large differences in the low temperature regime, the values of metallicity remain unaffected with $< \Delta [\text{Fe}/\text{H}] > = 0.00 \pm 0.00$ ($\sigma=0.02$) dex for the overall sample. The same effect appears for surface gravity with $< \Delta \log g > = -0.02 \pm 0.00$ ($\sigma=0.06$) dex, even though there is bigger scatter. This result suggests that surface gravity and metallicity are not

Table 3. Part of the derived stellar parameters of the 451 stars.

Star	T_{eff} (K)	$\log g$ (cm s^{-2})	ξ_t (km s^{-1})	$[Fe/H]$ (dex)	Mass (M_{\odot})	Age (Gyr)
...
HD 967	5595±17	4.59±0.02	0.90±0.05	-0.66±0.01	0.76±0.02	7.8±3.9
HD 1237	5489±39	4.46±0.11	1.04±0.06	0.06±0.03	0.92±0.02	1.8±1.7
HD 1320	5699±12	4.55±0.05	0.89±0.02	-0.26±0.01	0.91±0.02	2.4±2.0
HD 1388	5970±15	4.42±0.05	1.13±0.02	0.00±0.01	1.06±0.01	3.0±0.8
HD 1461	5740±16	4.36±0.03	0.94±0.02	0.18±0.01	1.05±0.02	2.0±1.1
HD 1581	5990±15	4.49±0.07	1.24±0.03	-0.18±0.01	1.03±0.02	2.2±1.0
HD 2025	4851±49	4.49±0.13	0.51±0.18	-0.37±0.02	0.71±0.01	4.6±4.0
HD 2071	5729±12	4.49±0.02	0.93±0.02	-0.08±0.01	0.99±0.01	0.9±0.6
HD 2638	5169±77	4.41±0.15	0.66±0.13	0.12±0.05	0.85±0.03	4.0±3.8
...

**Fig. 3.** The effect of temperature on the other parameters: metallicity (left panel) and surface gravity (right panel). The differences correspond to the values of this work minus the results of SO08.

as sensitive to the selection of the line list as temperature, for this temperature regime and for this method. The same effect appears for the cool stars as seen in Table 4.

The impact of the updated effective temperatures on $[Fe/H]$ and on $\log g$ is depicted on Fig. 3. The changes in metallicity show almost no correlation, within the errors, with the effective temperature. Only slightly higher metallicities appear for low temperatures, when comparing with SO08, yet within the errors. The surface gravities are also not correlated with temperature, even though there is high dispersion in the low temperature region. This result suggests that using this technique, a potential error in one of the parameters will not propagate to the others, avoiding systematic errors (see also Torres et al. 2012).

Metallicity has a key role in planet formation theories and is correlated with the planet frequency. The stellar sample, as mentioned before, is part of the HARPS GTO planet search program and it contains 102 up-to-date planet hosts. The metallicity distribution of the sample, presented in SO08, shows that the Jovian planets are preferentially found in metal-rich stars, in contrast to Neptune-like planets that do not seem to follow this trend, even though the number of these planets is small. The new metallicities derived with the new line list do not change this trend, making this correlation between stars and planets reliable even before adapting the new temperatures.

4. Ionization balance problem in cool stars.

In an LTE abundance analysis for solar type stars the ionization equilibrium should be satisfied. Using the standard spectroscopic method we force the abundance of Fe I and Fe II to agree. Surface gravity is determined from this tuning. However, there are many studies for cluster stars (Yong et al. 2004; Morel & Micela 2004; Schuler et al. 2006) and field stars (Allende Prieto et al. 2004; Ramírez et al. 2007, 2013) showing that the cooler dwarf stars deviate from the ionization balance, with systematic higher Fe II abundances over Fe I. The authors explain these discrepancies due to possible different scales in the stellar parameters (namely T_{eff} and $\log g$) or due to NLTE effects caused by the simplifications of the model atmospheres. The ionization balance of Fe I and Fe II can be investigated in this work by the behaviour of surface gravity. In fact, surface gravity mostly depends on Fe II lines, once the temperature scale is correct.

An essential test for the accuracy of surface gravity derived from spectroscopy is the comparison with surface gravity derived from parallaxes (trigonometric $\log g$), based on the fundamental relation:

$$\log \frac{g}{g_{\odot}} = \log \frac{M}{M_{\odot}} + 4 \log \frac{T_{\text{eff}}}{T_{\text{eff},\odot}} - \log \frac{L}{L_{\odot}}$$

We calculated the trigonometric $\log g$ using the new *Hipparcos* parallaxes (van Leeuwen 2007), V magnitudes, bolometric correction based on Flower (1996) and Torres (2010), solar magnitudes from (Bessell et al. 1998), the spectroscopic

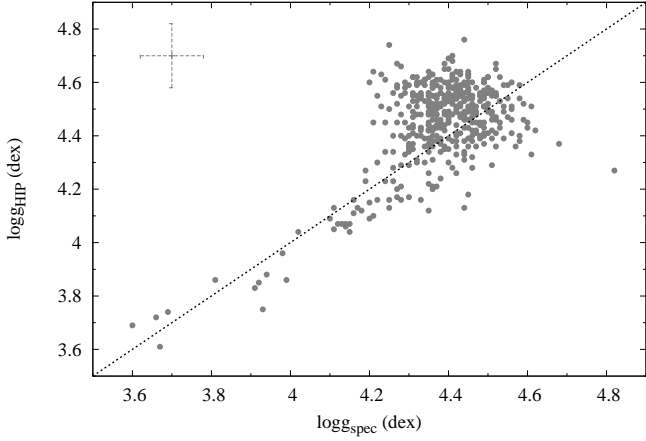


Fig. 4. Comparison of the surface gravities derived from spectroscopy, $\log g_{\text{spec}}$ of this work and from *Hipparcos* parallaxes, $\log g_{\text{HIP}}$. A mean error for both axes is given in the upper left part.

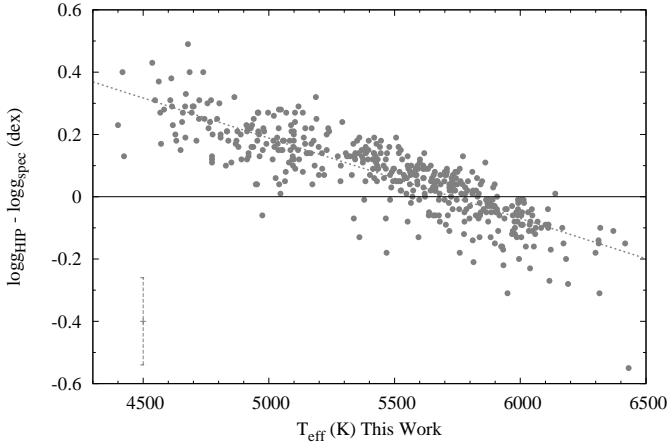


Fig. 5. Trigonometric $\log g_{\text{HIP}}$ minus $\log g_{\text{spec}}$ as a function of temperature. The dashed line represents a linear fit ($-2.578 \cdot 10^{-4} T_{\text{eff}} + 1.477$). The mean error is shown at the bottom left.

masses and T_{eff} . No correction for interstellar reddening is needed since all stars are less than 56 pc in distance.

In Fig. 4, we compare the spectroscopic $\log g_{\text{spec}}$ with the trigonometric $\log g_{\text{HIP}}$. At first glance, the spectroscopic $\log g$ agrees with the trigonometric ($\log g_{\text{HIP}} - \log g_{\text{spec}} > 0.07$ dex). However, a more careful look shows that there is a disagreement especially for the high values of $\log g$ (> 4.5 dex) where the spectroscopic gravities are underestimated, which is also observed in the work of SO08.

An interesting fact is that the differences between the trigonometric and spectroscopic $\log g$ are greater for the cooler stars (Fig. 5). There is a clear trend between the differences in $\log g$ and T_{eff} , where the underestimation of surface gravity in low temperatures becomes higher ($\log g_{\text{HIP}} - \log g_{\text{spec}} > 0.22$ dex for stars with $T_{\text{eff}} < 5000$ K). This is translated into systematically higher Fe II abundances over Fe I for low T_{eff} .

Such differences between Fe I and Fe II abundances are difficult to explain with model uncertainties and departures from LTE in the spectral line formation calculations as they are expected for the warmer stars of our sample where most iron is ionized (Lind et al. 2012). In addition, other model uncertainties related to granulation and activity of K-stars have been proposed

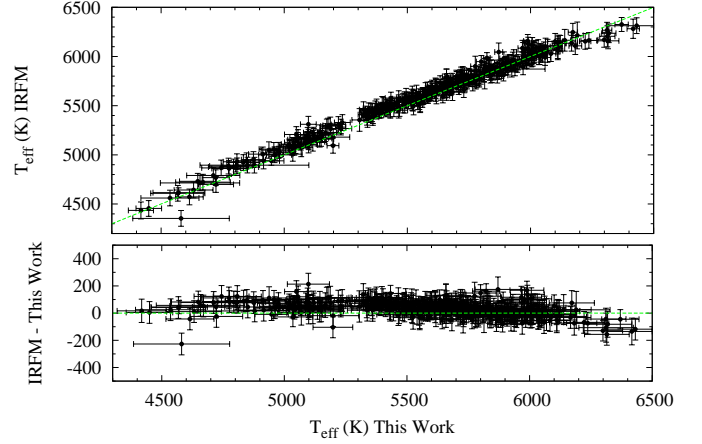


Fig. 6. Comparison between the temperatures derived from this work and the IRFM for stars in common.

to explain these differences even though these effects should be evident for young stars (Morel & Micela 2004; Schuler et al. 2010).

Other possible explanations for these differences have to do with the iron ionization method itself. We use Fe II lines that strongly depend on surface gravity. For solar-type stars, however, these lines are not sufficiently present leading to poorly constrains of $\log g$. On the other hand, the numerous Fe I lines are insensitive to $\log g$ changes.

We have to note though, that temperatures and metallicities derived using the ionization and excitation equilibrium of iron lines are shown to be mostly independent of the adopted surface gravity (Torres et al. 2012). Hence, the temperatures and metallicities derived with our spectroscopic method can be used as reference even if the derived spectroscopic surface gravities differ from the trigonometric values (see also Santos et al. 2013, submitted).

5. Comparison with other methods

In order to evaluate the consistency of our results, namely for T_{eff} , we compare them with other techniques. Here, we present a comparison with two different methods that are considered to be less model dependent, the Infrared Flux Method (IRFM) and interferometry, respectively.

5.1. The Infrared Flux Method - IRFM

The IRFM method (Blackwell & Shallis 1977) is a semi-direct method for determining stellar parameters. The principle of this method relies on the fact that the bolometric flux depends on the angular diameter and the effective temperature, as described by the Stefan-Boltzmann law, whereas the monochromatic flux in the infrared (IR) depends on the angular diameter but weakly on the effective temperature, this way the dependence on the angular diameter disappears:

$$\frac{f_{\text{bol}}}{f_{\text{IR}}} = \frac{\sigma T_{\text{eff}}^4}{f_{\text{IR}}(\text{model})}, \quad (1)$$

where f_{bol} is the measured bolometric flux, f_{IR} is the measured monochromatic IR flux and $f_{\text{IR}}(\text{model})$ is the monochromatic flux in the IR derived by the assuming model. The IRFM has the advantage that the dependence on the models is limited

Table 5. Comparison between the effective temperatures derived with different methods. σ represents the standard deviation and N the number of stars for the comparison.

Method	$\Delta T_{\text{eff}} (K)$	$\sigma (K)$	N
IRFM – this work	$+33 \pm 3$	54	347
IRFM – SO08	$+14 \pm 3$	61	347
IRFM – this work ($T_{\text{eff}} < 5000$ K)	$+42 \pm 12$	65	29
IRFM – SO08 ($T_{\text{eff}} < 5000$ K)	-86 ± 16	86	29
Interferometry – this work	-4 ± 33	98	9
Interferometry – SO08	-56 ± 35	105	9
Interferometry – IRFM	-98 ± 83	204	6

while the spectroscopic effective temperatures have considerable model dependence.

We compare our results with the work of Casagrande et al. (2010, 2011) that implement the IRFM method for a large sample of stars. The authors estimate the bolometric flux from multi-band photometric measurements in the optical $BV(RC)_C$ band and in the near-IR 2MASS JHK_S band. For the missing spectral regions, the flux is calculated by synthetic spectra computed from model atmospheres. The absolute calibration of Vega is based on its synthetic spectrum with an uncertainty of the zero point of ~ 15 K (Casagrande et al. 2010).

Figure 6 depicts the comparison between the spectroscopic temperatures and the IRFM for the stars in common. Temperatures of 341 stars are taken from Casagrande et al. (2011) using stars with direct application of the IRFM (*irfm* sample) and stars with T_{eff} derived from colour calibrations (*clbr* sample). Moreover, 6 stars were taken from Casagrande et al. (2010). The comparison between the results of this work and the IRFM shows good agreement for all temperature ranges. In particular, for the cooler temperature region, the differences in T_{eff} between this work and the IRFM are much smaller and more homogeneously distributed than between SO08 and the IRFM.

The mean differences in temperature for the comparison samples are shown in Table 5. It is clear that the differences in temperature for this work with the IRFM are constant throughout the temperature range, with a small offset of 33 K for the whole sample. For the cooler stars these differences are $\Delta T_{\text{eff}} = 42 \pm 12$ K that are much smaller than SO08 with $\Delta T_{\text{eff}} = -86 \pm 16$ K. Figure 7 shows the comparison of stars with $T_{\text{eff}} < 5000$ K. It is evident that the trend in ΔT_{eff} of SO08 mostly disappears with the new temperatures.

5.2. Interferometry

Precise measurements of stellar angular diameters are acquired through long baseline interferometry. The standard practice to determine the angular diameter is to fit the observed visibilities as a function of baseline to a uniform disk model. The angular size is connected to the more realistic limb darkened angular size (θ_{LD}) using correction factors from model atmospheres (Claret 2000). Temperature is then derived with the standard relation:

$$T_{\text{eff}} = \left(\frac{L}{4\pi\sigma R^2} \right)^{1/4} = \left(\frac{4f_{bol}}{\sigma\theta_{LD}^2} \right)^{1/4}, \quad (2)$$

where f_{bol} is usually calculated from the Spectral Energy Distribution.

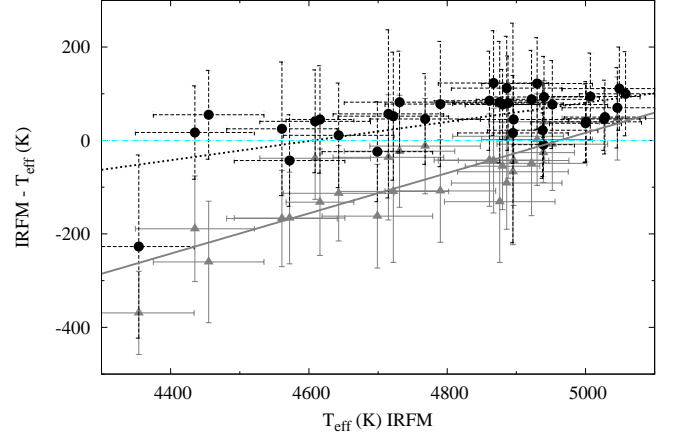


Fig. 7. Comparison between the difference in temperatures derived from IRFM - This Work (circles) and IRFM - SO08 (triangles). The dashed and the solid line depict the linear fits of the data with slopes: $+0.20 \pm 0.04$ and $+0.43 \pm 0.05$, respectively.

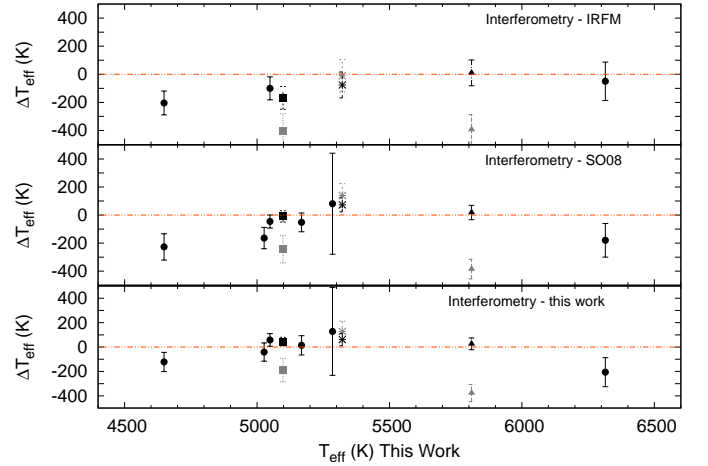


Fig. 8. Comparison between the spectroscopic, IRFM and the direct temperature measurements for stars in common with the literature. The x-axis corresponds to temperatures derived from this work. Square symbols are for the different temperatures for HD 26965, stars for HD 10700 and triangles for HD 146233. Stars with grey colour represent T_{eff} with high angular diameter uncertainty.

We compare our results with the temperatures derived from interferometry. Unfortunately, the number of stars with available angular diameters for this sample is only down to a few since these measurements are challenging for dwarfs due to their small photospheric disks that are difficult to resolve. We have 9 stars in common for the comparison with spectroscopy and 6 with the IRFM. We use only direct angular diameters and bolometric fluxes available in the literature from Table 6. Our results show better agreement for these stars to interferometry than when comparing with the values of SO08 and the IRFM (see Table 5). We have to note though, that the comparison sample is very small and the values of T_{eff} were derived from the *clbr* sample that is not the best representative for the IRFM precision.

The differences in temperatures for the different methods are plotted in Fig. 8. For 3 stars (HD 10700, HD 26965, HD 146233) we include angular diameter measurements from different authors that give different T_{eff} and are represented with different symbols. In the same figure, we see that from the stars with

Table 6. Interferometric data and derived temperatures for stars in common with our sample.

Star HD	θ_{LD} (mas)	$\Delta\theta/\theta$ (%)	$T_{\text{eff}}^{\text{Int}}$ (K)	$T_{\text{eff}}^{\text{SO08}}$ (K)	$T_{\text{eff}}^{\text{thiswork}}$ (K)	$T_{\text{eff}}^{\text{TRFM}}$ (K)	References
10700	2.022 ± 0.011	0.54	5383 ± 47	5310 ± 17	5322 ± 17	5459 ± 80	1, a
...	1.971 ± 0.050	2.54	5449 ± 83	2, a
11964	0.611 ± 0.081	13.25	5413 ± 359	5332 ± 22	5285 ± 21	-	3, b
19994	0.788 ± 0.026	3.30	6109 ± 111	6289 ± 46	6315 ± 44	6159 ± 80	3, b
22049	2.148 ± 0.029	1.35	5107 ± 21	5153 ± 42	5049 ± 48	5207 ± 80	4, c
23249	2.394 ± 0.029	1.21	4986 ± 57	5150 ± 51	5027 ± 48	-	5, a
26965	1.504 ± 0.006	0.40	5143 ± 14	5153 ± 38	5098 ± 32	5311 ± 80	6, d
...	1.650 ± 0.060	3.63	4910 ± 90	4, d
128621	6.001 ± 0.021	0.35	5182 ± 24	5234 ± 63	5168 ± 75	-	7, e
146233	0.676 ± 0.006	0.89	5836 ± 46	5818 ± 13	5810 ± 12	5826 ± 80	8, f
...	0.780 ± 0.017	2.18	5433 ± 69	9, f
209100	1.890 ± 0.020	1.06	4527 ± 29	4754 ± 89	4649 ± 73	4731 ± 80	4, g

Notes. References for θ_{LD} : (1) Teixeira et al. (2009); (2) Pijpers et al. (2003); (3) van Belle & von Braun (2009); (4) Kervella & Fouqué (2008); (5) Thévenin et al. (2005); (6) Boyajian et al. (2012b); (7) Kervella et al. (2003); (8) Bazot et al. (2011); (9) Boyajian et al. (2012a). References for f_{bol} : (a) Bruntt et al. (2010); (b) van Belle & von Braun (2009); (c) Cayrel et al. (2011); (d) Boyajian et al. (2012b); (e) Ramírez & Meléndez (2004); (f) Boyajian et al. (2012a); (g) Ramírez & Meléndez (2005).

multiple measurements, the ones with the smallest uncertainty ($\frac{\Delta\theta}{\theta}$ %) agree better with the temperatures derived with the spectroscopic and photometric methods. A precision better than 2% in the angular diameters corresponds to an accuracy of 1% in the effective temperatures, which is roughly 60 K at solar temperature, assuming no error in the bolometric flux. It is useful thus, to take the uncertainty of the angular diameter into consideration for a reliable determination of temperature.

6. New atmospheric parameters for cool planet hosts

In general, for planet host stars with effective temperature below 5200 K, the new set of parameters, derived using the line list presented in this paper, imply that the planet host stars have lower temperatures than previously published. This has implications for both their mass and radius determination. The lower temperatures imply lower stellar masses as well as lower stellar radii. This means that the derived planetary masses and radii (for transit planet cases) are also lower. As a consequence of the mass reduction, the semi-major axis of the orbits will also be smaller. The expected effects are however small, and no major revisions are expected to occur.

With this new line list, we re-derive the stellar parameters for 10 “cool” planet hosts already published in the literature and are not included in the 451 stellar sample of this work. We only consider GK dwarfs with an effective temperature lower than 5200 K whose planets were detected with the radial velocity technique from the CORALIE and HARPS GTO planet search samples. These planet hosts have been previously analyzed with high S/N spectra following the same procedure as this work but with different line lists.

In Table 7, the fundamental parameters based on the new line list are presented. The sixth column gives the reference of the previously published parameters. To explain the effect of these new parameters to mass more quantitative, we calculate the stellar masses from the Padova interface for both with the original and new parameters. We avoid to use the published stellar masses in order to compare uniformly. We find the maximum difference in mass to be 1.5% in absolute units which is negligible compared to the standard mass error.

7. [Cr I/Cr II] and [Ti I/Ti II] vs. T_{eff} with the new atmospheric parameters

The precise and accurate stellar parameters are, as well, very important for further analyzing stellar chemical abundances. The traditional spectroscopic abundance analysis methods require these parameters as input to compute the atmosphere models, hence the accuracy of the final elemental abundances depends on the accuracy of these input parameters. Different atoms and ions are not equally sensitive to all the stellar parameters. For example, ionized species are more sensitive to gravity variations than neutral species (e.g. Gilli et al. 2006; Neves et al. 2009; Adibekyan et al. 2012b).

Recently, Neves et al. (2009) and Adibekyan et al. (2012b) analyzing chemical abundances of the refractory elements of the HARPS sample stars, observed some unexpected trends with effective temperature. Particularly, they detected systematic trends of [X/H] or [X/Fe] with T_{eff} for some elements at low temperatures and found that [Cr I/Cr II] and [Ti I/Ti II] abundance ratios gradually increase with decreasing effective temperature when $T_{\text{eff}} \lesssim 5000$ K. Similar trends for different elements with T_{eff} have been already noted in the literature (see e.g. Valenti & Fischer 2005; Preston et al. 2006; Gilli et al. 2006; Lai et al. 2008; Suda et al. 2011). Different explanations of the mentioned trends are discussed in the literature. The unexpected trends in the low temperature regime may be due to the stronger line blending and may also be connected to either deviations from excitation or ionization equilibrium, or to problems associated with the differential analysis (Neves et al. 2009). A possible explanation for the observed trends with T_{eff} could also be an incorrect T - τ relationship in the adopted model atmospheres (Lai et al. 2008) or NLTE effects (Bodaghee et al. 2003). Summarizing, it can be assumed that the observed trends are probably not an effect of stellar evolution, and uncertainties in atmospheric models are the dominant effect in measurements (see also the discussion in Adibekyan et al. 2012b).

In Fig. 9, we plot the [Cr I/Cr II] and [Ti I/Ti II] abundance ratios derived using the stellar parameters of SO08 and this work as a function of the T_{eff} for stars cooler than 5000 K. This plot is useful to ensure that the ionization equilibrium enforced on the Fe II lines is acceptable to other elements. As can be seen the slopes of the abundance ratios with new parameters are very gentle. The new slope of [Cr I/Cr II] per 1000 K is

Table 7. Updated stellar parameters for previously analyzed planet hosts.

Name	T_{eff} (K)	$\log g$ (cm s^{-2})	[Fe/H] (dex)	ξ_t (km s^{-1})	Reference
BD-082823	4648 ± 135	4.33 ± 0.32	0.00 ± 0.08	0.27 ± 0.81	1
HD 3651	5182 ± 79	4.30 ± 0.16	0.12 ± 0.05	0.66 ± 0.15	2
HD 13445	5114 ± 61	4.55 ± 0.13	-0.29 ± 0.04	0.66 ± 0.15	2
HD 20868	4720 ± 91	4.24 ± 0.22	0.08 ± 0.06	0.47 ± 0.31	1
HD 99492	4815 ± 184	4.28 ± 0.46	0.24 ± 0.12	0.50 ± 0.56	4
HD 125595	4596 ± 235	4.25 ± 0.63	0.10 ± 0.14	0.14 ± 1.41	3
HD 128311	4778 ± 75	4.35 ± 0.17	-0.03 ± 0.02	0.82 ± 0.16	2
HD 192263	4906 ± 57	4.36 ± 0.17	-0.07 ± 0.02	0.78 ± 0.12	2
HD 215497	5003 ± 103	4.26 ± 0.26	0.25 ± 0.05	0.61 ± 0.22	1
HIP 5158	4673 ± 175	4.24 ± 0.47	0.22 ± 0.12	0.34 ± 1.09	1

References. (1) Sousa et al. (2011); (2) Santos et al. (2004); (3) Ségransan et al. (2011); (4) Santos et al. (2005).

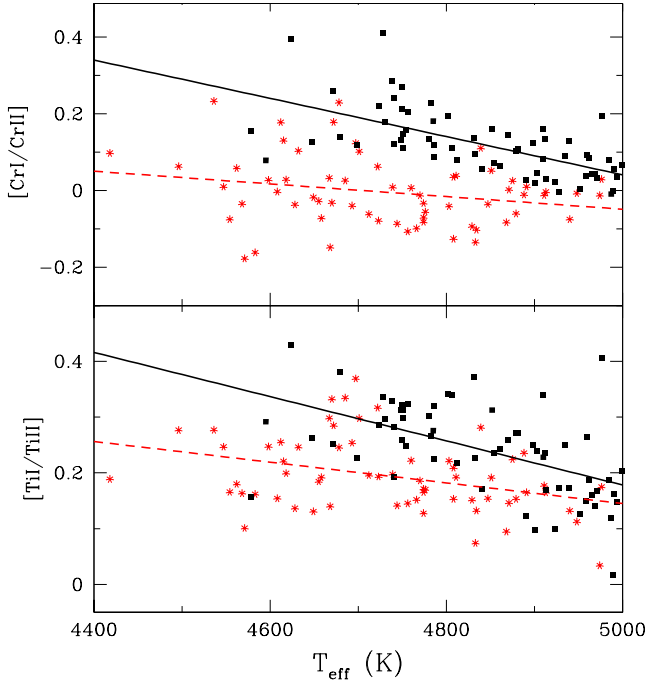


Fig. 9. $[\text{Cr I/Cr II}]$ and $[\text{Ti I/Ti II}]$ as a function of effective temperature. The black squares and red asterisks correspond to the abundance ratios derived using old and new stellar parameters, respectively. The solid and dashed lines depict the linear fits of the data.

-0.16 ± 0.08 , whereas the slope with the parameters of SO08 is -0.49 ± 0.08 . The new slope of $[\text{Ti I/Ti II}]$ is also improved a lot and is -0.18 ± 0.06 dex per 1000 K. For comparison the slope of SO08 is -0.39 ± 0.08 dex. Although the trend with T_{eff} is weak, there is a shift of about 0.2 dex for the $[\text{Ti I/Ti II}]$ ratio. This shift is difficult to connect to the still possible uncertainties in the stellar parameters and its exact nature still remains to be clarified. Probably, one (or more) of the above mentioned effects can be responsible for that. Unfortunately, in the literature there is no available NLTE calculations for the Ti I and Ti II lines used in our study, and it is difficult to estimate the NLTE effect for the $[\text{Ti I/Ti II}]$ ratio.

Summarizing, this independent test shows that the new stellar parameters derived from the iron lines more carefully chosen for cooler stars make the observed $[\text{Cr I/Cr II}]$ and $[\text{Ti I/Ti II}]$ trends with T_{eff} much weaker.

8. Conclusions

In this work, we present a new iron line list in order to correct for discrepancies in temperatures for a sample of 451 stars that is part of the HARPS GTO program using the cool star HD 21749 as a reference. The quality of the line list plays a key role for temperature determination especially for the K type stars. The new line list is compiled in order to eliminate blended lines in the spectra of these stars that suffer more from such effects. We also apply a limitation to very strong and very weak lines that usually cause errors in the EW measurements (see Sect. 2.1).

We derived the stellar parameters for the 451 stars of the sample in a homogeneous way with the new line list and compare our results with the work of SO08, where the authors followed the same analysis but with an expanded line list.

We find very good agreement for the high and intermediate ranges in temperature with SO08. The differences appear in the lower temperatures below ~ 5000 K. In addition, surface gravity and metallicity remain unaffected with the new line list. These results are very important since accurate parameters of planet host stars are essential to characterize the planets. In particular, the metallicities of this sample are used for the investigation between the stars and their planet hosts. The fact that the new metallicities are not different from the old ones supports the already established correlation between the planet host stars and planet frequency.

The agreement with the IRFM that is considered to be a less model dependent technique, suggests that our parameters are more precise at least for the cooler stars compared to the previous study of SO08 where the existing trend in low temperatures between the spectroscopic and the photometric temperatures disappears with the new line list. In addition, we compare our new temperatures with the ones derived with interferometry for 9 star with available angular diameter measurements. Even though the sample is very small, a comparison between these results shows very good agreement. Attention should be made to temperatures derived with the angular diameters that have larger uncertainties.

Having checked the reliability of the new line list, we apply it to determine fundamental parameters for some of the planet hosts with low temperatures, which were previously analyzed with different line lists with available high S/N spectra. Finally, we use the new parameters for stars with $T_{\text{eff}} < 5000$ K to significantly reduce the trends of $[\text{Cr I/Cr II}]$ and $[\text{Ti I/Ti II}]$ abundance ratios with temperature.

Acknowledgements. The authors acknowledge the referee, I. Ramírez, for his comments that helped improve this paper. This work was supported by the European Research Council/European Community under the FP7 through Starting Grant agreement number 239953. N.C.S. also acknowledges the support

from Fundação para a Ciência e a Tecnologia (FCT) through program Ciência 2007 funded by FCT/MCTES (Portugal) and POPH/FSE (EC), and in the form of grant reference PTDC/CTE-AST/098528/2008. V.Zh.A. and S.G.S. are supported by grants SFRH/BPD/70574/2010 and SFRH/BPD/47611/2008, respectively.

References

- Adibekyan, V. Z., Delgado Mena, E., Sousa, S. G., et al. 2012a, *A&A*, 547, A36
 Adibekyan, V. Z., Sousa, S. G., Santos, N. C., et al. 2012b, *A&A*, 545, A32
 Albrecht, S., Winn, J. N., Johnson, J. A., et al. 2012, *ApJ*, 757, 18
 Allende Prieto, C., Barklem, P. S., Lambert, D. L., & Cunha, K. 2004, *A&A*, 420, 183
 Ammler-von Eiff, M., Santos, N. C., Sousa, S. G., et al. 2009, *A&A*, 507, 523
 Bazot, M., Ireland, M. J., Huber, D., et al. 2011, *A&A*, 526, L4
 Bergemann, M., Lind, K., Collet, R., Magic, Z., & Asplund, M. 2012, *MNRAS*, 427, 27
 Bessell, M. S., Castelli, F., & Plez, B. 1998, *A&A*, 333, 231
 Blackwell, D. E. & Shallis, M. J. 1977, *MNRAS*, 180, 177
 Bodaghee, A., Santos, N. C., Israelian, G., & Mayor, M. 2003, *A&A*, 404, 715
 Bouchy, F., Hébrard, G., Udry, S., et al. 2009, *A&A*, 505, 853
 Boyajian, T. S., McAlister, H. A., van Belle, G., et al. 2012a, *ApJ*, 746, 101
 Boyajian, T. S., von Braun, K., van Belle, G., et al. 2012b, *ApJ*, 757, 112
 Bruntt, H., Bedding, T. R., Quirion, P.-O., et al. 2010, *MNRAS*, 405, 1907
 Buchhave, L. A., Latham, D. W., Johansen, A., et al. 2012, *Nature*, 486, 375
 Casagrande, L., Portinari, L., & Flynn, C. 2006, *MNRAS*, 373, 13
 Casagrande, L., Ramírez, I., Meléndez, J., Bessell, M., & Asplund, M. 2010, *A&A*, 512, A54
 Casagrande, L., Schönrich, R., Asplund, M., et al. 2011, *A&A*, 530, A138
 Cayrel, R., van't Veer-Menneret, C., Allard, N. F., & Stehlé, C. 2011, *A&A*, 531, A83
 Claret, A. 2000, *A&A*, 363, 1081
 Flower, P. J. 1996, *ApJ*, 469, 355
 Fuhrmann, K. 2004, *Astronomische Nachrichten*, 325, 3
 Gilli, G., Israelian, G., Ecuivillon, A., Santos, N. C., & Mayor, M. 2006, *A&A*, 449, 723
 Girardi, L., Bressan, A., Bertelli, G., & Chiosi, C. 2000, *A&AS*, 141, 371
 Gray, D. F. 1994, *PASP*, 106, 1248
 Guillot, T., Santos, N. C., Pont, F., et al. 2006, *A&A*, 453, L21
 Kervella, P. & Fouqué, P. 2008, *A&A*, 491, 855
 Kervella, P., Thévenin, F., Ségransan, D., et al. 2003, *A&A*, 404, 1087
 Kovtyukh, V. V., Soubiran, C., Belik, S. I., & Gorlova, N. I. 2003, *A&A*, 411, 559
 Kurucz, R. 1993, *ATLAS9 Stellar Atmosphere Programs and 2 km/s grid*. Kurucz CD-ROM No. 13. Cambridge, Mass.: Smithsonian Astrophysical Observatory, 1993., 13
 Lai, D. K., Bolte, M., Johnson, J. A., et al. 2008, *ApJ*, 681, 1524
 Lind, K., Bergemann, M., & Asplund, M. 2012, *MNRAS*, 427, 50
 Lo Curto, G., Mayor, M., Benz, W., et al. 2010, *A&A*, 512, A48
 Mashonkina, L., Gehren, T., Shi, J.-R., Korn, A. J., & Grupp, F. 2011, *A&A*, 528, A87
 Mayor, M., Pepe, F., Queloz, D., et al. 2003, *The Messenger*, 114, 20
 Morel, T. & Micela, G. 2004, *A&A*, 423, 677
 Neves, V., Santos, N. C., Sousa, S. G., Correia, A. C. M., & Israelian, G. 2009, *A&A*, 497, 563
 Nissen, P. E. 1981, *A&A*, 97, 145
 Pepe, F., Lovis, C., Ségransan, D., et al. 2011, *A&A*, 534, A58
 Pijpers, F. P., Teixeira, T. C., Garcia, P. J., et al. 2003, *A&A*, 406, L15
 Press, W. H., Teukolsky, S. A., Vetterling, W. T., & Flannery, B. P. 1992, *Numerical recipes in FORTRAN. The art of scientific computing*
 Preston, G. W., Sneden, C., Thompson, I. B., Shectman, S. A., & Burley, G. S. 2006, *AJ*, 132, 85
 Ramírez, I., Allende Prieto, C., & Lambert, D. L. 2007, *A&A*, 465, 271
 Ramírez, I., Allende Prieto, C., & Lambert, D. L. 2013, *ApJ*, 764, 78
 Ramírez, I. & Meléndez, J. 2004, *ApJ*, 609, 417
 Ramírez, I. & Meléndez, J. 2005, *ApJ*, 626, 446
 Reddy, B. E., Tomkin, J., Lambert, D. L., & Allende Prieto, C. 2003, *MNRAS*, 340, 304
 Ruchti, G. R., Bergemann, M., Serenelli, A., Casagrande, L., & Lind, K. 2013, *MNRAS*, 429, 126
 Santos, N. C., Israelian, G., & Mayor, M. 2004, *A&A*, 415, 1153
 Santos, N. C., Israelian, G., Mayor, M., et al. 2005, *A&A*, 437, 1127
 Santos, N. C., Sousa, S. G., Mortier, A., et al. 2013, submitted
 Schuler, S. C., King, J. R., Terndrup, D. M., et al. 2006, *ApJ*, 636, 432
 Schuler, S. C., Plunkett, A. L., King, J. R., & Pinsonneault, M. H. 2010, *PASP*, 122, 766
 Ségransan, D., Mayor, M., Udry, S., et al. 2011, *A&A*, 535, A54

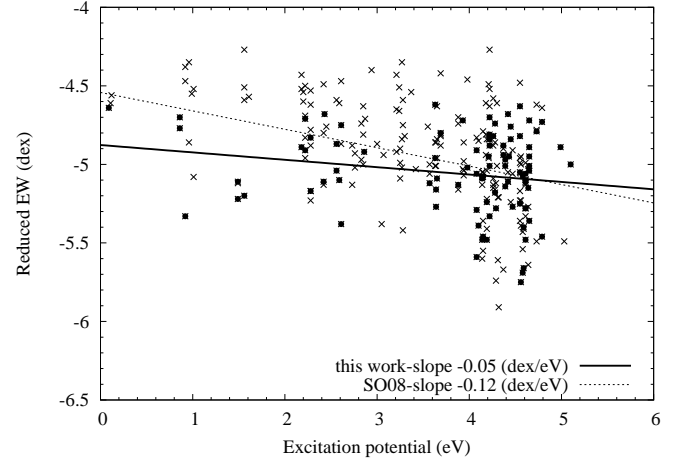


Fig. A.1. Reduced EW versus excitation potential for the line list of this work (circles) and for SO08 (crosses). The slope of the linear fits is also depicted.

- Sneden, C. A. 1973, PhD thesis, THE UNIVERSITY OF TEXAS AT AUSTIN.
 Sousa, S. G., Santos, N. C., Israelian, G., Mayor, M., & Monteiro, M. J. P. F. G. 2007, *A&A*, 469, 783
 Sousa, S. G., Santos, N. C., Israelian, G., Mayor, M., & Udry, S. 2011, *A&A*, 533, A141
 Sousa, S. G., Santos, N. C., Mayor, M., et al. 2008, *A&A*, 487, 373
 Sozzetti, A. 2004, *MNRAS*, 354, 1194
 Suda, T., Yamada, S., Katsuta, Y., et al. 2011, *MNRAS*, 412, 843
 Teixeira, T. C., Kjeldsen, H., Bedding, T. R., et al. 2009, *A&A*, 494, 237
 Thévenin, F., Kervella, P., Pichon, B., et al. 2005, *A&A*, 436, 253
 Torres, G. 2010, *AJ*, 140, 1158
 Torres, G., Fischer, D. A., Sozzetti, A., et al. 2012, *ApJ*, 757, 161
 Torres, G., Winn, J. N., & Holman, M. J. 2008, *ApJ*, 677, 1324
 Udry, S., Bonfils, X., Delfosse, X., et al. 2007, *A&A*, 469, L43
 Udry, S., Mayor, M., Naef, D., et al. 2000, *A&A*, 356, 590
 Valenti, J. A. & Fischer, D. A. 2005, *ApJS*, 159, 141
 van Belle, G. T. & von Braun, K. 2009, *ApJ*, 694, 1085
 van Leeuwen, F. 2007, *A&A*, 474, 653
 Vogt, S. S., Marcy, G. W., Butler, R. P., & Apps, K. 2000, *ApJ*, 536, 902
 Winn, J. N., Fabrycky, D., Albrecht, S., & Johnson, J. A. 2010, *ApJ*, 718, L145
 Yong, D., Lambert, D. L., Allende Prieto, C., & Paulson, D. B. 2004, *ApJ*, 603, 697

Appendix A: Reduced EW versus Excitation Potential

We show the correlation between the reduced EW and excitation potential (Fig. A.1). For the line list of this work there is much smaller correlation than in the work of SO08 suggesting small degeneracies between ξ_i and T_{eff} .

Appendix B: The microturbulence relationship

Microturbulence is taken into consideration for abundance analyses to reconcile differences between the observed and predicted from models equivalent widths of strong lines. Previous studies of FGK dwarfs have shown that ξ_i depends on T_{eff} and $\log g$ (e.g. Nissen 1981; Reddy et al. 2003; Allende Prieto et al. 2004; Adibekyan et al. 2012a; Ramírez et al. 2013). Using a linear regression analysis to the new parameters of the sample, we derive the following expression:

$$\xi_i = 6.932 \times 10^{-4} T_{\text{eff}} - 0.348 \log g - 1.437. \quad (\text{B.1})$$

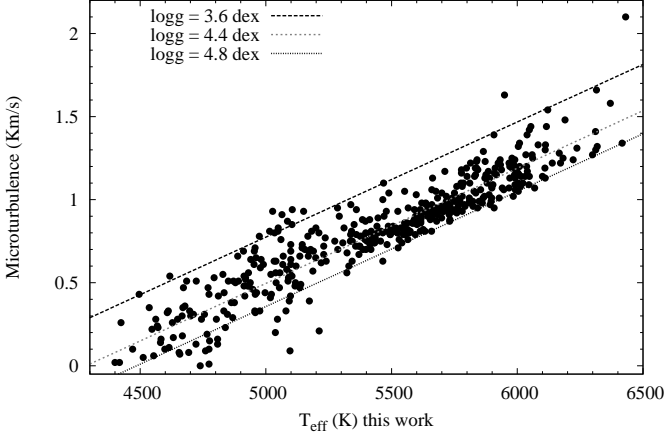


Fig. B.1. Correlation of microturbulence with temperature and surface gravity as described by equation B.1. We set $\log g$ to the highest, average and lowest values of the sample.

Here, ξ_t is in km/s, T_{eff} and $\log g$ are in their traditional units. The parameters of the stars in the sample range:

$$\begin{aligned} 4400 < T_{\text{eff}} < 6400 \text{ K}, \\ 3.6 < \log g < 4.8 \text{ dex}, \\ -0.8 < [\text{Fe}/\text{H}] < 0.4 \text{ dex}. \end{aligned}$$

The new derived parameters indicate a linear dependence on temperature for a set value of surface gravity. In Fig. B.1, we see the dependence of microturbulence on temperature for a set of $\log g$ values. Microturbulence clearly increases with temperature and decreases with surface gravity.